Geometric Mechanics

Homework 13

Due on January 11

1. Show that if we introduce a **cosmological constant** $\Lambda \in \mathbb{R}$ in the Einstein equation,

$$Ric = 4\pi\rho(2\nu\otimes\nu + g) + \Lambda g$$

then the equations for the FLRW models become

$$\begin{cases} \frac{\dot{a}^2}{2} - \frac{\alpha}{a} - \frac{\Lambda}{6}a^2 = -\frac{k}{2} \\ \frac{4\pi}{3}a^3\rho = \alpha \end{cases}$$

Analyze the possible behaviors of the function a(t). (Remark: It is currently thought that there exists indeed a positive cosmological constant, also known as **dark energy**. The model favored by experimental observations seems to be $k=0, \Lambda>0$).

- 2. Consider two galaxies in a FLRW model, whose spatial locations can be assumed to be r=0 and $(r,\theta,\varphi)=(r_1,\theta_1,\varphi_1)$.
 - (a) Show that the family (reparameterized) null geodesics connecting the first galaxy to the second galaxy can be written as

$$(t, r, \theta, \varphi) = (t(r, t_0), r, \theta_1, \varphi_1)$$
 $(0 < r < r_1),$

where $t(r, t_0)$ is the solution of

$$\begin{cases} \frac{dt}{dr} = \frac{a(t)}{\sqrt{1 - kr^2}} \\ t(0, t_0) = t_0 \end{cases}$$

- (b) Prove that $\frac{\partial t}{\partial t_0}(r_1,t_0)=\frac{a(t_1)}{a(t_0)}$, where $t_1=t(r_1,t_0)$.
- (c) The **redshift** of the light propagating from the first galaxy to the second galaxy is defined as

$$z = \frac{\partial t}{\partial t_0}(r_1, t_0) - 1 = \frac{a(t_1)}{a(t_0)} - 1.$$

This light is spread over a sphere of radius $R=a(t_1)r_1$, and so its brightness is inversely proportional to \mathbb{R}^2 . Compute R as a function of z for the following FLRW models:

- (i) Milne universe $(k=-1, \alpha=\Lambda=0)$, for which a(t)=t;
- (ii) Flat de Sitter universe ($k=\alpha=0$, $\Lambda=3H^2$), for which $a(t)=e^{Ht}$;
- (iii) Einstein-de Sitter universe ($k=\Lambda=0,~\alpha=2/9t_1{}^2$), for which $a(t)=(t/t_1)^{2/3}$.

(Remark: The brightness of distant galaxies is further reduced by a factor of $(1+z)^2$, since each photon has frequency, hence energy, (1+z) times smaller at reception, and the rate of detection of photons is (1+z) times smaller than the rate of emission; with this correction, R can be deduced from the observed brightness for galaxies of known luminosity, and the correct FLRW model can be chosen as the one whose curve R = R(z) best fits observations).